

Shock Structure and Shock Heating in the Galactic Central Molecular Zone

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Abstract. We present maps of a large number of dense molecular gas tracers across the Central Molecular Zone of our Galaxy. The data were taken with the CSIRO/CASS Mopra telescope in Large Projects in the 1.3 cm, 7 mm, and 3 mm wavelength regime. Here, we focus on the brightness of the shock tracers SiO and HNC, molecules that are liberated from dust grains under strong (SiO) and weak (HNC) shocks. The shocks may have occurred when the gas enters the bar regions and the shock differences could be due to differences in the moving cloud mass. Based on tracers of ionizing photons, it is unlikely that the morphological differences are due to selective photo-dissociation of the molecules. We also observe direct heating of molecular gas in strongly shocked zones, with a high SiO/HNC ratios, where temperatures are determined from the transitions of ammonia. Strong shocks appear to be the most efficient heating source of molecular gas, apart from high energy emission emitted by the central supermassive black hole Sgr A* and the processes within the extreme star formation region Sgr B2.

Keywords. ISM: molecules — Galaxy: center — radio lines: ISM

1. Introduction

The Central Molecular Zone (CMZ) covers the inner ~ 500 pc of the Milky Way and contains a substantial amount of molecular gas (several $10^7 M_\odot$ e.g. Oka et al. 1998). Models show that the gas may be funneled along the central Galactic bar into the CMZ and eventually accretes on x_2 orbits that cover the inner ~ 200 pc (e.g. Morris & Serabyn 1996). Eventually the gas will be destroyed by environmental influences, during the process of star formation, or, in rare cases, the dense gas will be transported all the way further toward Sgr A*, the central supermassive black hole of the Milky Way.

In this paper, we focus on material that is formed during shocks. HNC, for example, is a molecule that is formed on dust grains and liberated into the gas phase by weak shocks (Meier & Turner 2005). Stronger shocks can be traced by SiO where the silicon is provided by shock destruction of the dust grains themselves (Martin-Pintado et al. 1997). Since the ISM conditions in the Galactic Center are dominated by shock physics, we use these two molecules to trace the environment as the gas accumulates in the CMZ. Jones et al. (2012) presents maps of both molecules using the ATNF/CASS Mopra† telescope and Ott et al. (2008) shows single dish Mopra maps of ammonia, a molecule which is used as a thermometer to trace shock heating.

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2. The Shock Structure of Inflowing Gas

In Fig. 1 we show the distribution of the HNC0 ($4_{04} - 3_{03}$) and SiO (2-1) line emission in the CMZ. In particular the regions east of Sgr B2 show a distinct pattern of alternating strong SiO and HNC0 lines. This hints to alternating stronger and weaker shocks, where the shocks might have occurred at the time the gas entered the large Galactic bar or along x_1 orbits as the gas moves toward the Galactic Center. Eventually the gas may accrete on the 100 pc ring (likely the location of the x_2 orbits) that was discovered in *HERSCHEL* dust maps (Molinari et al. 2011). This scenario is in agreement with gas heating by shocks. At the positions where the SiO emission dominates, the ammonia temperature map shows much higher temperatures, compared to the places with HNC0 dominating. In addition, the velocity dispersion increases at the same morphological features. The shock heated gas is some of the hottest gas in the Galactic Center, outside of the environments of compact sources.

References

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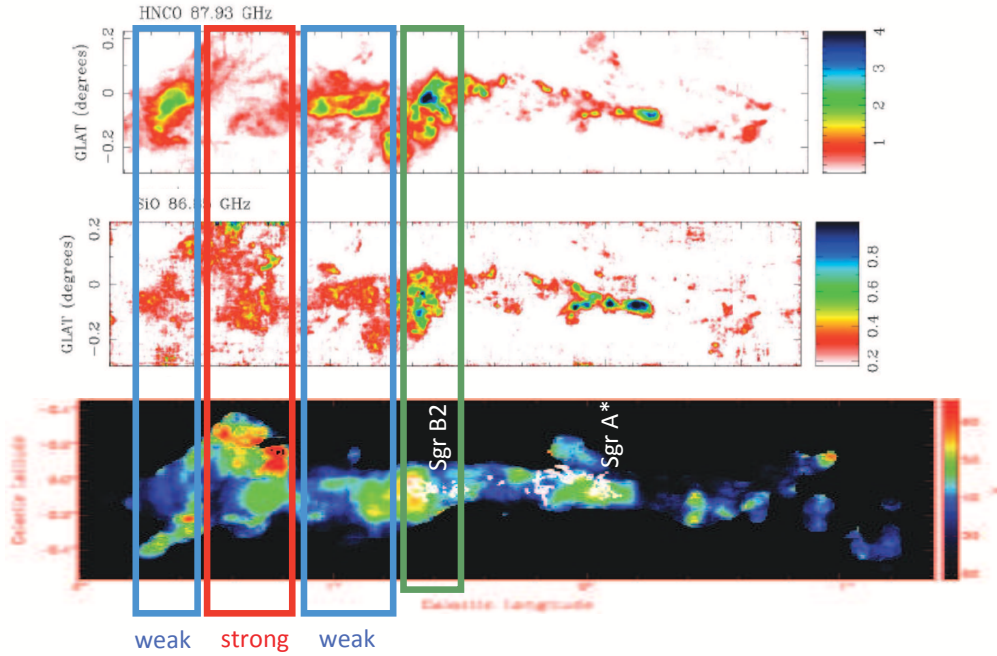


Figure 1. Shock tracers in the CMZ: *Top:* HNC0, *Middle:* SiO. At the bottom we show a gas temperature map as derived from NH₃. The regions with weak shocks are marked in *blue*, the strongly shocked regions in *red*. The *green* box centers on the star forming region Sgr B2 and marks the likely point where the x_1 orbits accrete on the 100 pc ring, the location of the x_2 orbits (color figure available in the online version).